



# Direct measurement of $S(E)$ factor for the ${}^9\text{Be}(p,a/d)$ reactions in low energy region (18-100 keV)

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# Motivation

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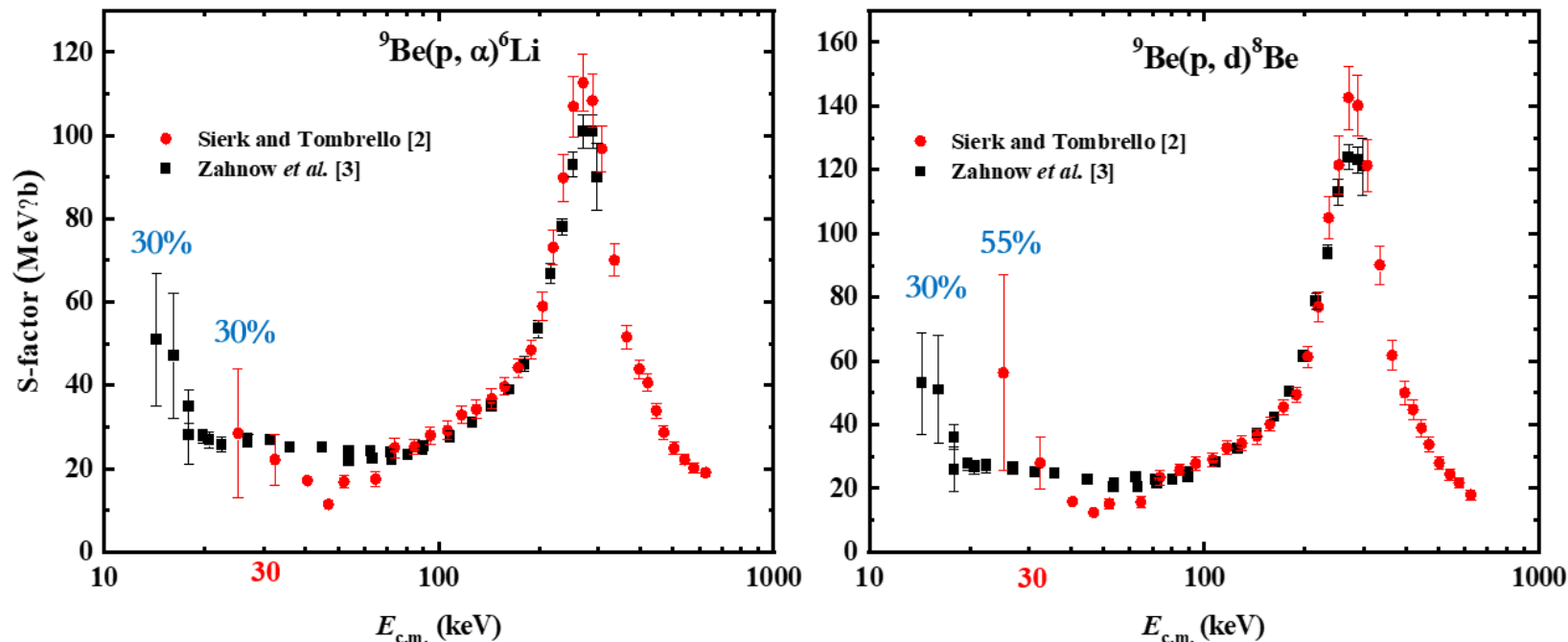
## Why reaction rate of the $p$ - ${}^9\text{Be}$ reactions measure in low energy?

- Be depletion mechanisms must be well understood to differentiate between cosmic-ray and big-bang production models.
- In both stellar and primordial environments, LiBeB are mainly destroyed by proton-capture reactions via the  $(p,\alpha)$  channel with a Gamow energy  $E_G$  ranging from 10 keV (for stellar nucleosynthesis) to 100 keV (for primordial nucleosynthesis).
- Primordial beryllium abundance could reveal insights into the Big Bang. Beryllium's reaction rate could be useful for constraining the nonstandard BBN models together with the Li and B.

# Introduction

## p-<sup>9</sup>Be reactions:

### Direct measurement



The astrophysical  $S(E)$  of the  ${}^9\text{Be}(p, \alpha){}^6\text{Li}$  (left) and  ${}^9\text{Be}(p, d){}^8\text{Be}$  (right) reactions

Large errors in low energy region.

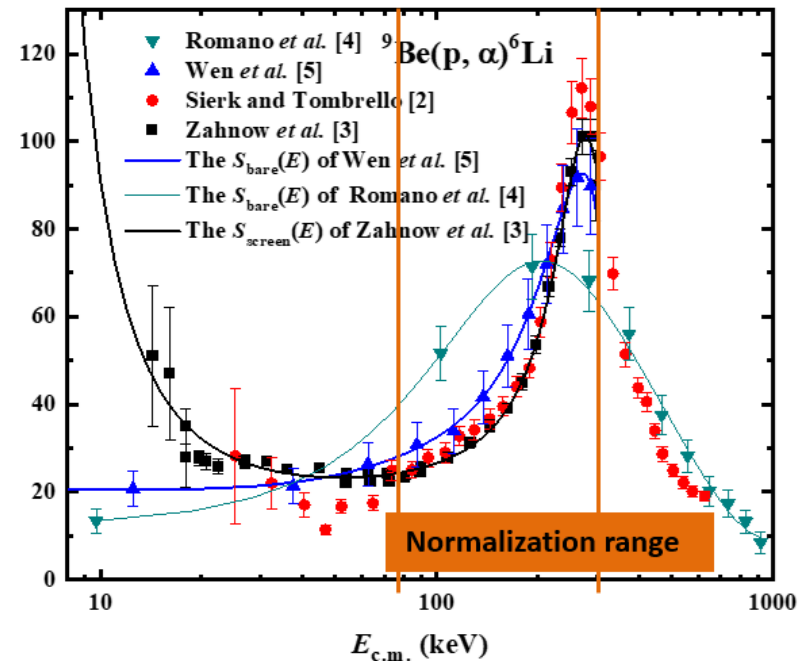
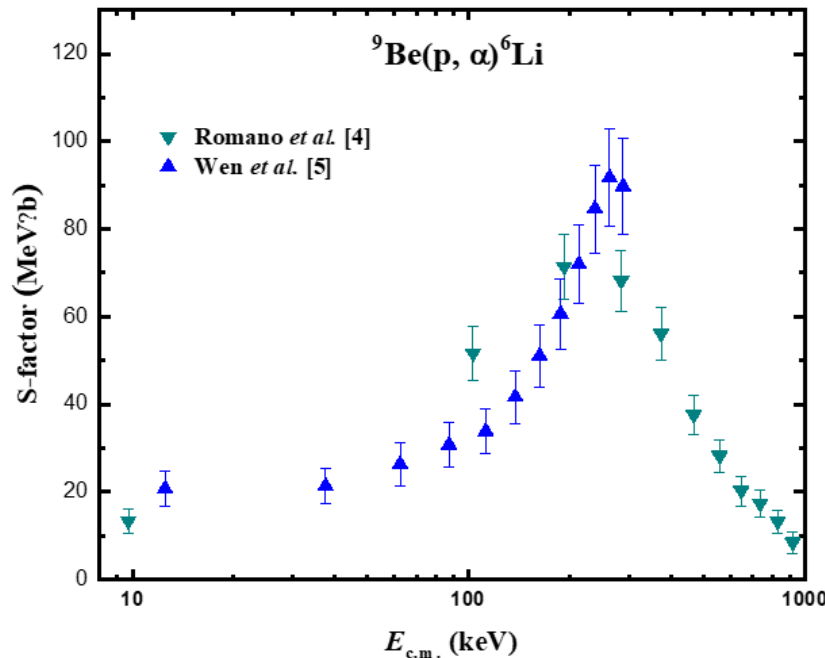
[2]. A.J. Sierk, T.A. Tombrello, The  ${}^9\text{Be}(p, \alpha)$  and  $(p, d)$  cross sections at low energies, Nuclear Physics A, 210 (1973) 341-354.

[3]. D. Zahnnow, et al., Low-energy  $S(E)$  factor of  ${}^9\text{Be}(p, \alpha){}^6\text{Li}$  and  ${}^9\text{Be}(p, d){}^8\text{Be}$ , Nuclear Physics A, 359 (1997) 211-218.

# Introduction

## p-<sup>9</sup>Be reactions:

### Trojan horse method (THM)



The astrophysical  $S(E)$  of the  ${}^9\text{Be}(p, \alpha){}^6\text{Li}$  (left) and  ${}^9\text{Be}(p, d){}^8\text{Be}$  (right) reactions

larger resonance width compared to direct measurements.

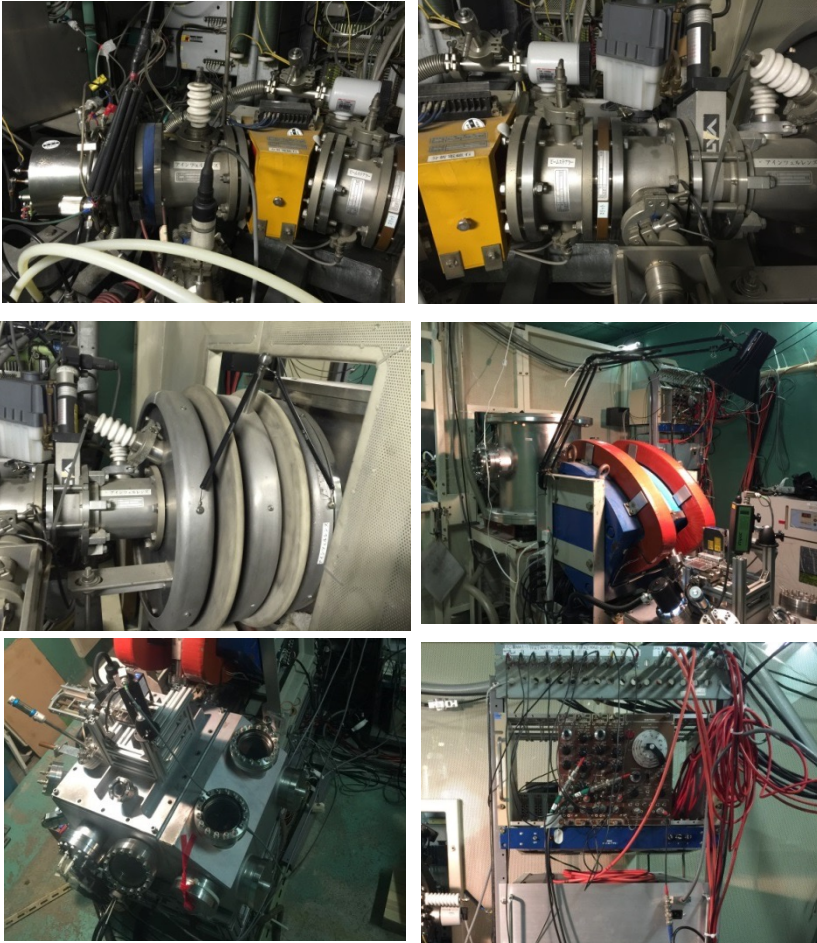
➤ Measure the cross section in the energy range from 18 to 100 keV.

[4]. S. Romano, et al., Study of the  ${}^9\text{Be}(p, \alpha){}^6\text{Li}$  reaction via the Trojan Horse Method, The European Physical Journal A, 27 (2006) 221-225.

[5]. Q.G. Wen, et al., Trojan horse method applied to  ${}^9\text{Be}(p, \alpha){}^6\text{Li}$  at astrophysical energies, Phys.rev.c, 78 (2008).

# Experimental setup

## Low-energy high-intensity accelerator



1-100 keV high-current accelerator, Tohoku University

### Accelerator parameters

Accelerator:  $E_{p/d}$ : 2 ~ 100 keV

$I_{\max} = 500 \mu\text{A}$

$\Delta E = \pm 35 \text{ eV}$

Filament life ~ 150 hours

Vacuum: ion source ~  $10^{-3} \text{ Pa}$

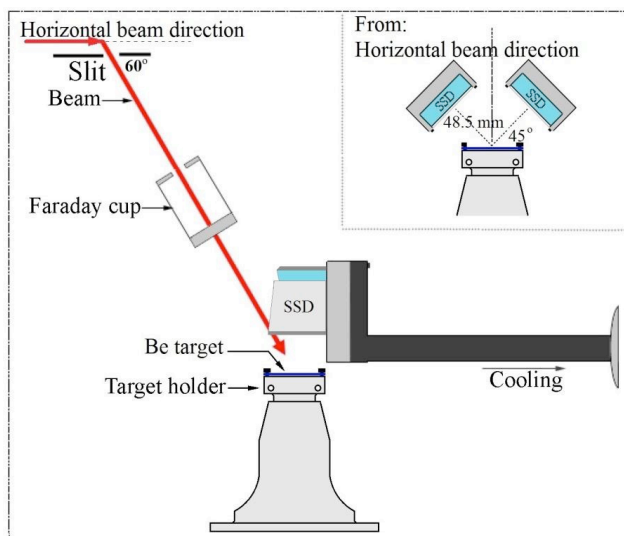
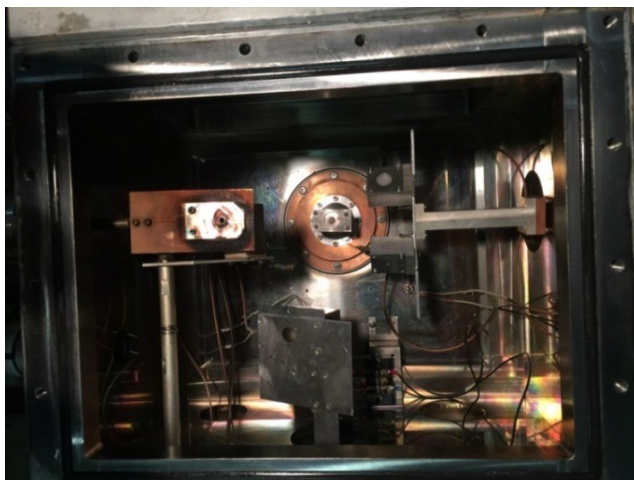
reaction chamber ~  $10^{-5} \text{ Pa}$

Monitor: Temperature-Infrared thermometer

Beam—Faraday cup (F.C.)

Cooling: water ( $5^\circ\text{C}$ )

# Experimental setup



The configuration of SSD detectors in chamber

## Experimental conditions

Beams:  $H^+$ :  $28 \leq E_{lab} \leq 100$  keV

$H_3^+$ :  $54 \leq E_{lab} \leq 102$  keV ( $\Delta E=2$  keV/amu)

$I_A$ : 5~130  $\mu A$

Spot size: 8 mm

Target: Be (99%),  $0.1 \times 25 \times 25$  mm<sup>3</sup>

SSD: Area: 450 mm<sup>2</sup>

Thickness: 300  $\mu m$

Al foil: 1.0  $\mu m$

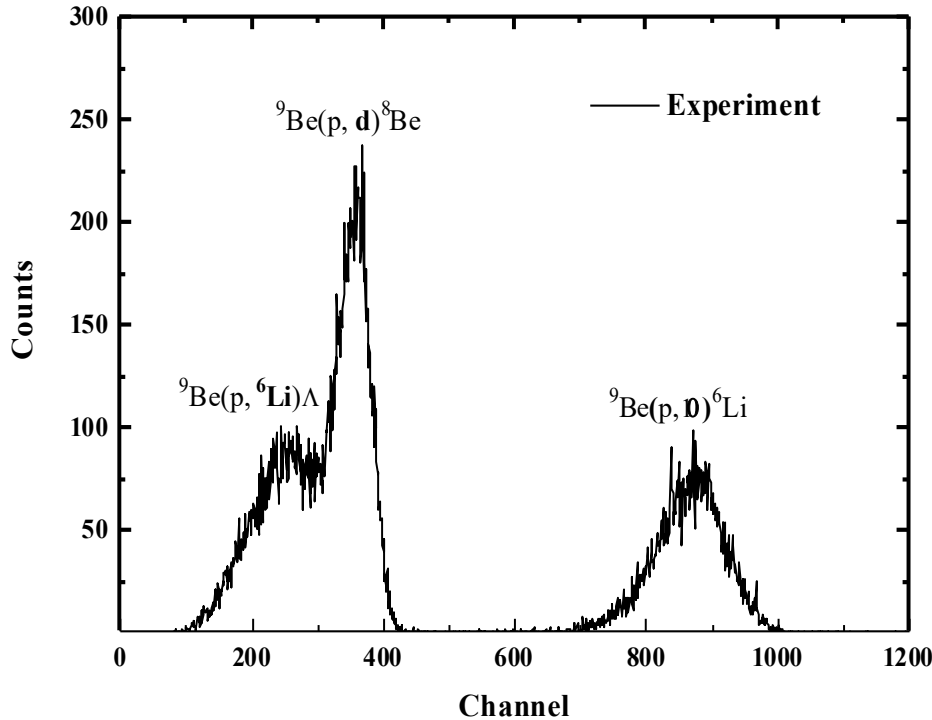
Detection angles: 127°

Solid angle:  $0.42 \pm 0.01$  sr

**Method: Thick target yield**

# Results and analysis

## Experimental spectrum



The spectrum of p-<sup>9</sup>Be reactions at  $E_p = 100$  keV

The thick-target yield:

$$Yield(E_i) = \frac{Counts}{N_p}$$

$$Yield(E_p) = \frac{N_p N_t \Delta\Omega_{lab}}{4\pi} \int_0^{E_p} \frac{d\Omega_{c.m.}}{d\Omega_{lab}} W(\theta, E) \sigma(E) \times \left(\frac{dE}{dx}\right)^{-1} dE$$

$$\sigma_{bare}(E) = \frac{S_{bare}(E)}{E} \exp(-2\pi\eta(E))$$

$$Yield_d = Yield_{total} - Yield_{Li}$$

$$S(E_{eff}) = \frac{Y_{exp}(E_0) - Y_{exp}(E_0 - \Delta)}{\frac{N_p N_t \Delta\Omega_{lab}}{4\pi} \int_{E_0 - \Delta}^{E_0} \frac{d\Omega_{c.m.}}{d\Omega_{lab}} W(\theta, E) \frac{\exp(-2\pi\eta)}{E} \left(\frac{dE}{dx}\right)^{-1} dE}$$

$$E_{eff} = E_0 - \Delta E + \Delta E \left\{ -\frac{\sigma_2}{\sigma_1 - \sigma_2} + \left[ \frac{\sigma_1^2 + \sigma_2^2}{2(\sigma_1 - \sigma_2)^2} \right]^{1/2} \right\}$$



*Thank you!*

*Welcome to my poster board!*

